

Le vide quantique autour des étoiles à neutrons

Carlo RIZZO

Laboratoire National des Champs Magnétiques intenses

Université de Toulouse

et

CNRS

UNIVERSITÉ
PAUL
SABATIER



TOULOUSE III



Electromagnetic energy density

$$U = \vec{E} \frac{\partial L}{\partial \vec{E}} - L$$

The vacuum is Lorentz and CPT invariant :

$$L = \frac{1}{2} F + aF^2 + bG^2 + \dots \quad \text{where} \quad F = \left(\epsilon_0 E^2 - \frac{B^2}{\mu_0} \right); \quad G = \sqrt{\frac{\epsilon_0}{\mu_0}} (\vec{E} \cdot \vec{B})$$

↑
Maxwell

$$\vec{D} = \epsilon_0 \vec{E} + \vec{P} = \frac{\partial L}{\partial \vec{E}}$$

$$\vec{H} = \frac{\vec{B}}{\mu_0} - \vec{M} = -\frac{\partial L}{\partial \vec{B}}$$

Constitutive
équations

1935-1936 Kochel, Euler, Heisenberg

$$\frac{b}{a} = 7$$

$$a = \frac{2}{45} \frac{\alpha^2 \hbar^3}{m_e^4 c^5}$$

H.Euler et K.Kochel, *Naturwiss.* **23** (1935); W.Heisenberg et H.Euler, *Z. Phys.* **38** (1936) 714

$$\vec{P} = 4a\epsilon_0 \vec{E}F + 14a \sqrt{\frac{\epsilon_0}{\mu_0}} \vec{B}G$$

$$\vec{M} = -4a \frac{\vec{B}}{\mu_0} F + 14a \sqrt{\frac{\epsilon_0}{\mu_0}} \vec{E}G$$

The vacuum is a non linear optical medium !

Vacuum in an external field

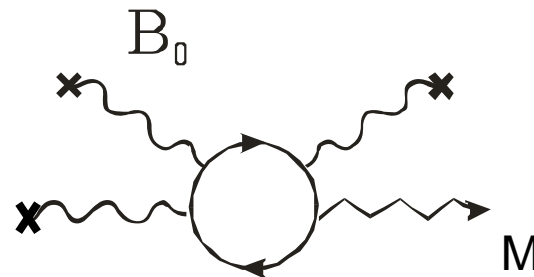
$$\vec{B}_0$$

$$\vec{P} = 0$$

$$\vec{M} = 4a \frac{B_0^2}{\mu_0^2} \vec{B}_0 = \frac{2}{45} \alpha \left(\frac{B_0^2}{B_c^2} \right) \frac{\vec{B}_0}{\mu_0}$$

where

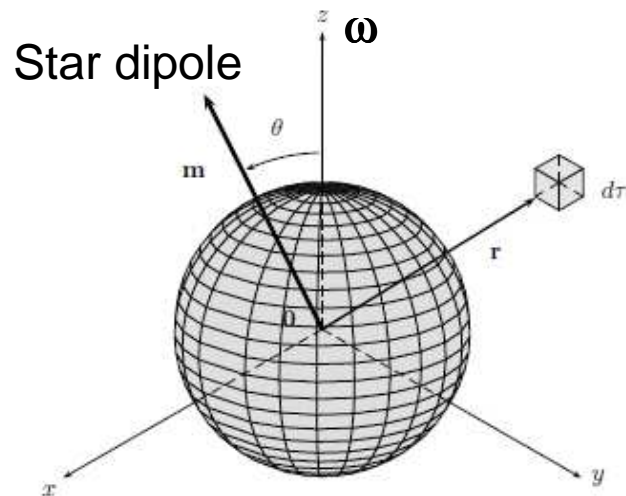
$$B_c = m_e^2 c^3 / e \hbar \simeq 4.4 \cdot 10^9 \text{ T}$$



Vacuum can be magnetized !

Neutron stars

M=Solar mass Star radius = 10 km

 $B_0 > B_c$?! Magnetars

$$\mathbf{B}(\mathbf{r}, t) \simeq \left(\frac{\mu_0}{4\pi} \right) \left[\frac{3\mathbf{r}(\mathbf{m}(t - r/c) \cdot \mathbf{r})}{r^5} - \frac{\mathbf{m}(t - r/c)}{r^3} \right]$$

$$d\mathbf{m}_{qv}(\mathbf{r}, t) = \frac{\alpha B^2(\mathbf{r}, t)}{2\pi B_c^2 \mu_0} f_{qv}(B^2(\mathbf{r}, t)) \mathbf{B}(\mathbf{r}, t) d\tau$$

$$d\dot{\mathbf{E}}_{qv} = -(\mathbf{m}(t + r/c) \times d\mathbf{B}_{qv}(\mathbf{0}, t + r/c)) \boldsymbol{\omega} \cdot \mathbf{u}_z$$

Quantum Vacuum Friction

$$\dot{\mathbf{E}}_{qv} \simeq \alpha \left(\frac{18\pi^2}{45} \right) \frac{\sin^2 \theta}{B_c^2 \mu_0 c} \frac{B_0^4 R^4}{p^2}$$

Experimental data are : $P, \dot{P}, \ddot{P}, \dddot{P}$ ← rarely
 often

Is QVF in
 contadiction with
 existing data ?

$$\dot{E} = 4\pi^2 I \dot{P} P^{-3}$$

/ inertia moment assumed as constant in time

Classical radiation from rotating magnetic dipole :

$$\dot{E}_r = \left(\frac{128\pi^5}{3} \right) \frac{\sin^2 \theta}{\mu_0 c^3} \frac{B_0^2 R^6}{p^4}$$

B_0 is given assuming that $\dot{E} = 4\pi^2 I \dot{P} P^{-3} = \dot{E}_r = \left(\frac{128\pi^5}{3} \right) \frac{\sin^2 \theta}{\mu_0 c^3} \frac{B_0^2 R^6}{p^4}$

$$B_0 \propto \sqrt{P \dot{P}}$$

and assuming $\sin \theta = 1$

The braking factor :

$$n = \frac{\nu \ddot{\nu}}{\dot{\nu}^2} \quad n = 2 - \frac{P \ddot{P}}{\dot{P}^2} \quad \nu = 1/P$$

Classical dipole radiation

$$\dot{\nu} = -\frac{\dot{E}_r}{4\pi^2 I \nu} = -K_r \nu^3 \quad \text{and} \quad n_{cl}=3 \quad !?$$

The problem is that no neutron star shows an $n = 3$!

Thus

$$\dot{E} = 4\pi^2 I \dot{P} P^{-3} \neq \dot{E}_r = \left(\frac{128\pi^5}{3} \right) \frac{\sin^2 \theta}{\mu_0 c^3} \frac{B_0^2 R^6}{p^4}$$

!

B_0 ??

Nobody cares ?!

QVF

$$\dot{P} = \left(\frac{32\pi^3}{3I} \right) \frac{B_0^2 R^6}{\mu_0 c^3 P} + \alpha \left(\frac{3}{16I} \right) \frac{B_0^4 R^4 P}{B_c^2 \mu_0 c}$$

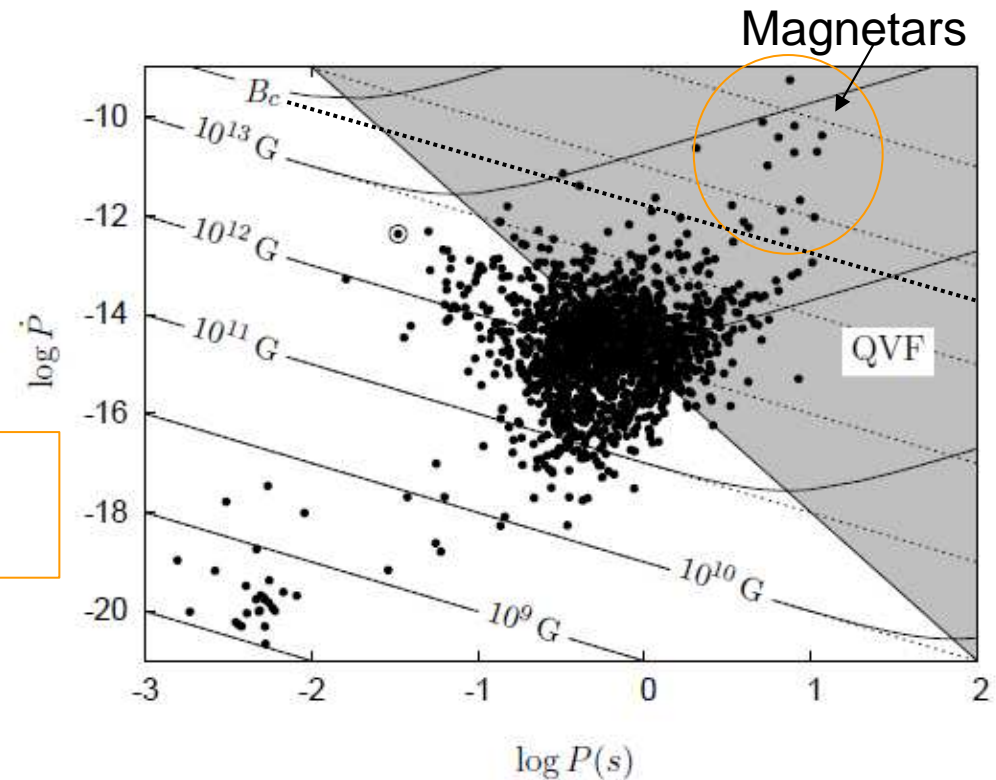


FIG. 1: $P - \dot{P}$ -diagram for the currently known pulsars. Pulsars located in the upper right area (gray) have an energy loss process dominated by QVF. The Crab pulsar is marked as a circle filled with a dot.

« classical » values of B_0
 « QVF » values of B_0 _____

$$n = \frac{\nu \ddot{\nu}}{\dot{\nu}^2}$$

$$n = 2 - \frac{P\ddot{P}}{\dot{P}^2}$$

What about n ?

QVF can reproduce the measured value of n .

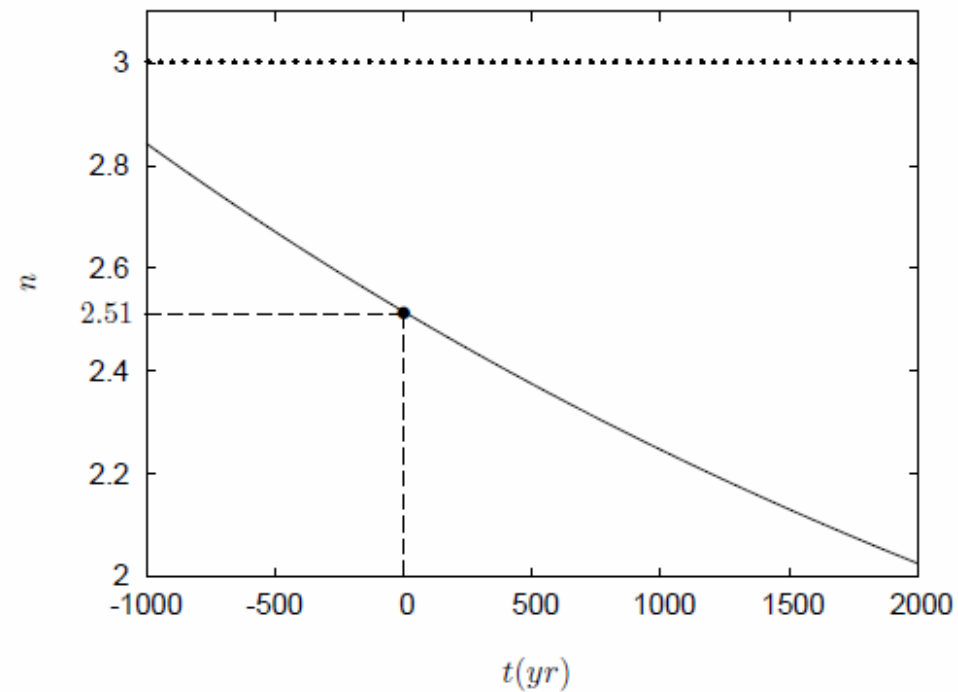


FIG. 3: Braking index of the Crab pulsar as a function of time when QVF is taken into account (full line) and when QVF is neglected (dotted line). $t = 0$ corresponds to current time.

QVF can be used to predict $\dot{n} = f(P, \dot{P}, \ddot{P}, \ddot{\ddot{P}})$

For the **Crab pulsar** we get

$-1.8421 \cdot 10^{-11} \text{ s}^{-1}$ vs $1.5676 \cdot 10^{-13} \text{ s}^{-1}$,

whereas

for the **B1509-58 pulsar** we get

$-1.3550 \cdot 10^{-11} \text{ s}^{-1}$ vs $-1.2500 \cdot 10^{-11} \text{ s}^{-1}$!!!.

The only two stars in the ATNF catalogue for which $\ddot{\ddot{P}}$ has been measured.

There are no contradictions between existing experimental data and QVF !

Does QVF exist or not ?

We need more experimental data !

Observations of QVF in pulsars would confirm the existence of one of the very few macroscopic effects of quantum vacuum,

while

if QVF is invalidate this would indicate for the first time that the usual QED model of the quantum vacuum used to calculate the QVF effect needs to be mended.

Both issues are very important !

Le vide quantique autour des étoiles à neutrons

D. Bakalov, G.F. Bignami, A. Dupays, C. Rizzo

arXiv:0804.1841[astro-ph] published EPL

arXiv:1010.0597[astro-ph] submitted PRD

UNIVERSITÉ
PAUL
SABATIER



TOULOUSE III

carlo.rizzo@incmi.cnrs.fr

